

Snowmass2021 - Letter of Interest

Astrophysical signatures of the QCD axion and axion-like particles

Thematic Areas: (check all that apply /■)

- (CF1) Dark Matter: Particle Like
- (CF2) Dark Matter: Wavelike
- (CF3) Dark Matter: Cosmic Probes
- (CF4) Dark Energy and Cosmic Acceleration: The Modern Universe
- (CF5) Dark Energy and Cosmic Acceleration: Cosmic Dawn and Before
- (CF6) Dark Energy and Cosmic Acceleration: Complementarity of Probes and New Facilities
- (CF7) Cosmic Probes of Fundamental Physics
- (Other) TF9

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Abstract:

We describe a program making use of astrophysical data sets to search for evidence of the Quantum Chromodynamics axion and axion-like particles. This LOI focuses on four related physical phenomena: (i) observable modifications to stellar cooling and core-collapse supernovae due to axions; (ii) hard X -ray spectra produced from evolved stars in the presence of axions from axion-photon conversion in astrophysical magnetic fields; (iii) spectral modulations induced by low-mass axions in X -ray and γ -ray data sets from astrophysical sources; (iv) nearly monochromatic radio lines produced by axion dark matter from either axion-photon conversion in NS magnetospheres or stimulated decay in the ambient Galaxy. Searches for the phenomena described above have already set some of the strongest constraints to-date on the axion across a wide range of possible masses and couplings, and we outline future theoretical and observational efforts needed to further improve sensitivity and possibly lead to a discovery.

The Quantum Chromodynamics (QCD) axion is a well-motivated extension to the Standard Model that may solve the strong- CP problem of the neutron electric dipole moment and also explain the observed dark matter (DM)^{1–7}. The axion may arise naturally in String Theory constructions, which predict a large number of axion-like particles in the low-energy spectrum^{8–10}. The QCD axion a is defined through its coupling to QCD: $\mathcal{L} = -(g^2/32\pi^2 f_a)aG\tilde{G}$, with f_a the axion decay constant, G the QCD field strength, and g the QCD coupling constant. Below the QCD confinement scale this operator generates a mass m_a for the axion: $m_a \approx 10 \mu\text{eV}(10^{12} \text{ GeV}/f_a)$. The axion couples also to photons and matter; in the infrared, these couplings can be described by

$$\mathcal{L} \supset \frac{C_{a\gamma\gamma}\alpha_{\text{EM}}}{2\pi f_a} \cdot \frac{1}{4}aF^{\mu\nu}\tilde{F}_{\mu\nu} + \sum_f \frac{C_{af}}{2f_a} \cdot (\partial_\mu a)\bar{f}\gamma^\mu\gamma_5 f, \quad (1)$$

with $C_{a\gamma\gamma}$ and C_{af} dimensionless coupling constants and f denoting the quark (or meson and baryon at low energies) and lepton fields. Note that it is common to define the axion-photon coupling $g_{a\gamma\gamma} \equiv C_{a\gamma\gamma} \cdot \alpha_{\text{EM}}/(2\pi f_a)$ and the dimensionless matter couplings $g_{aff} \equiv C_{af}m_f/f_a$, with m_f the fermion mass. Axion-like particles couple to photons and matter through the terms in (1) but do not (necessarily) couple to QCD; thus, their masses m_a do not need to have a simple relation to their couplings. Both the QCD axion and axion-like particles may make up the observed DM abundance over a wide range of possible masses and couplings through the mis-alignment mechanism and other non-thermal production mechanisms^{11,12}.

The axion parameter space is currently the subject of significant laboratory effort¹³. Some of the leading running and planned experiments are ADMX^{14–16}, HAYSTAC¹⁷, CAPP^{18,19}, ABRACADABRA^{20–22}, DM Radio^{23–26}, CASPER²⁷, MADMAX²⁸, CAST^{29–32}, IAXO³³, and ALPS-II³⁴. However, indirect searches for axions with astrophysical observations play an important and complementary role to the laboratory searches. Here we discuss astrophysical observations that span a range of wavelengths and physical phenomena, including: signatures of axions in stellar cooling and core-collapse supernovae; X -ray searches for axion-photon conversion from axions produced in stars; X -ray and γ -ray signatures of axion-induced spectral modulations from astrophysical sources; and radio telescope searches for axion dark matter. The current constraints on the axion-photon coupling $g_{a\gamma\gamma}$ and the axion-electron coupling g_{aee} as functions of the axion mass m_a are summarized in Fig. 1. We note that the strongest constraints on these couplings and the other axion-matter couplings, across a broad range of possible masses, come from astrophysical observations. In this Letter of Interest (LOI) we argue that these search strategies are promising directions for future focus that could lead to the discovery of axions.

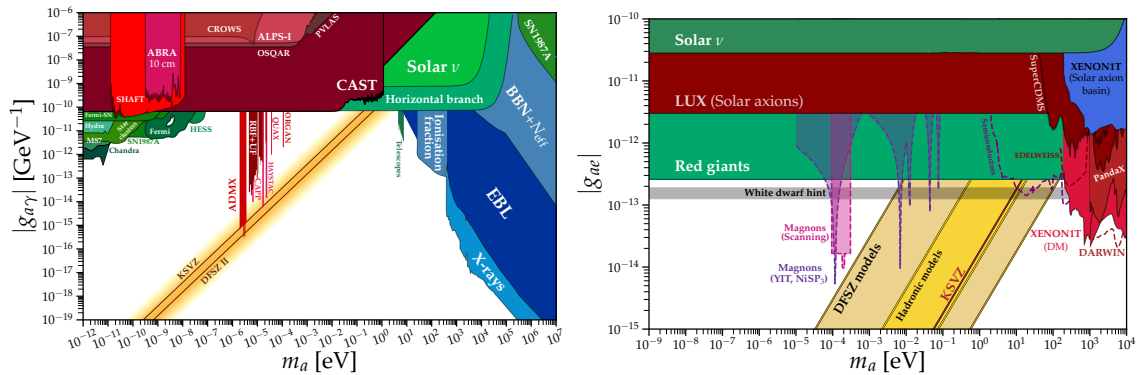


Figure 1: Current constraints on the axion-photon and axion-electron couplings (other axion-matter couplings not shown)³⁵. The astrophysical searches discussed in this LOI are the most sensitive probes of axions over large regions the possible parameter space. Improving upon the sensitivity of these searches, through theoretical and observational developments, is crucial to maximize the chances of discovering axions.

Axion effects on stars and supernovae.—Axion production in stars has long been understood to provide a pathway by which stars may cool. Some of the strongest constraints on the axion-photon and the axion-matter couplings arise from stellar cooling observations, such as horizontal branch star, white dwarf (WD) star, neutron star (NS), and red giant star cooling (see Fig. 1)³⁶. It is also known that axion emission from the core of a hot proto-NS might impact the observable supernova neutrino breakout burst and cooling phase^{37–46}. Prospects for improving the constraints on axion emission in the course of a supernova explosion could allow for improvement in current limits⁴⁷ or could plausibly lead to a detection of new physics⁴⁸, and an analysis that takes into account upcoming experimental sensitivities of next-generation neutrino experiments⁴⁹ is important to pursue. Likewise, the study of the effects of axions on low-mass stellar populations have returned both positive and negative results^{36,43,50–59}, and investigating the effects of observational and modeling uncertainties on these inferences is a critical next step for understanding how axions impact stars. Finally, the implications of axion emission for other phases of stellar evolution and different stellar populations provide promising new signals, such as the impending detection of the black hole mass gap^{60,61} in Advanced LIGO/Virgo.

X-ray searches for stellar axions.—Stellar axions may be directly observable as X-rays after they escape the stars and convert to photons in either the stellar or galactic magnetic fields^{62–67}. Strong constraints on $g_{a\gamma\gamma}$ for low-mass axions have been set using NuSTAR data from Super Star Clusters⁶⁸ (see Fig. 1). Observations with *XMM-Newton* and *Chandra* towards nearby NSs⁶⁶ and WDs⁶⁵ have set the strongest constraints to-date on $g_{a\gamma\gamma} \times g_{an}$ (in addition to a possible detection⁶⁶) and $g_{a\gamma\gamma} \times g_{ae}$, respectively, where g_{an} is the axion-neutron coupling. Significant progress across the X-ray axion searches could be made by future X-ray telescopes such as ATHENA, and better understanding the theoretical uncertainties both on the axion production in, *e.g.*, NSs and on the conversion probabilities in the stellar and galactic magnetic fields would also be fruitful.

X-ray and γ -ray spectral modulations.—The interconversion of photons and axions in astrophysical magnetic fields can lead to energy-dependent modulations in the spectra of γ - and X-ray sources. Searches for such an effect have been performed in γ -rays using data from, *e.g.*, *Fermi-LAT* and H.E.S.S towards Galactic and extragalactic sources, such as Galactic pulsars⁶⁹, supernova remnants^{70,71}, and active galaxies^{72–77}, yielding a possible detection. X-ray modulation searches with *Chandra* towards *e.g.* NGC 1275⁷⁸ and M87⁷⁹ have set the strongest constraints on axion with masses below $\sim 10^{-12}$ eV (see Fig. 1), consistently with previous studies^{80–83}. These analyses require modeling of the conversion in the intra-cluster, extragalactic and Galactic magnetic fields and are thus subject to uncertainties on these fields; better understanding these uncertainties is crucial for future progress. Next-generation instruments such as ATHENA and X-ray polarimeters are expected to further increase the sensitivity of these searches to axions^{84,85}.

Radio searches for axion dark matter.—Axion DM may fall into the gravitational potential wells of NSs and then resonantly convert to nearly monochromatic radio signatures in the strong magnetic fields in the magnetospheres^{86–92}. NS magnetospheres give the photon an effective mass that increases monotonically as one approaches the surface; level crossing and resonant conversion occurs at the radius where the plasma mass matches the axion mass. Dedicated searches for the axion-induced radio signal were performed in^{92–94}; no evidence for axions were found but strong constraints on $g_{a\gamma\gamma}$ were set for $m_a \sim 10 \mu\text{eV}$. In recent years, the Breakthrough Listen project⁹⁵ on the Green Bank Telescope has made publicly available⁹⁶ a large set of target spectra in the 1–12 GHz range. This opens up the exciting possibility for model-independent stimulated axion decay searches^{90,97,98} and NS searches across a wide range of masses. The projected reach of a dedicated search towards *e.g.* the GC with the planned Square Kilometer Array (SKA) is sufficient to cover larger regions of QCD axion DM parameter space for masses $m_a \sim 10 \mu\text{eV}$ ^{89,90}. More work is needed to (i) understand the theoretical uncertainties related to calculating the axion-induced radio flux, and (ii) developing search strategies for existing and future observatories.

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